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" 2. That the first day of the Egyptian year corresponded, throughout the series, either to about August 29, or August 30, of the Julian calendar, but with an uncertainty as respects the exact day; the Egyptian calendar, however, being probably identical with the fixed Alexandrian calendar.

" 3. That year 9 of the first series commenced in August, 105, A.D., making the first series of Egyptian years correspond to the termination of the reign of the emperor Trajan, as was conjectured by M. Brugsch. Consequently:—

Egyptian.					Julian.		
Series I.	Year 9	1st month,	1st day	=	A.D. 105	August 29	or 30
—	— 19	—	—	=	115	—	—
Series II.	Year 1	—	—	=	116	—	—
—	— 17	—	—	=	132	—	—

" 4. That the Egyptian places are either calculated places, or places found by some equivalent method; and are not the result of actual observation.

" With respect to the accuracy of the places calculated for comparison with the Egyptian places, it may be mentioned that in calculating the heliocentric longitudes, the equation of the centre was always applied: for *Jupiter* and *Saturn* the great equation was also applied: no others were taken into account. In finding the geocentric longitudes, the true radius vector was always used, excepting for some of the places of *Venus* first calculated, for which the mean distances of the Earth and *Venus* were used."

Occultation of Antares. Observed by Capt. Shadwell, R.N.

	Portsmouth M.T. h m s
Disappearance (dark limb)	11 23 24·2
Reappearance (bright limb)	12 41 29·4

"At the disappearance the weather was quite clear, and the phenomenon was instantaneous.

"At the reappearance a slight haze was passing over the moon, but the time is probably true within a second.

"*R. N. College, Portsmouth, June 16th,* 1856."

Occultation of Antares, June 16th, 1856.

Observed by T. W. Burr, Esq.

" Notwithstanding that the night was generally unfavourable for observation—the sky being mostly covered with hazy clouds

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—a small space round the moon became clear prior to the occultation, which was therefore very well seen.

	Sidereal Time.
	h m s
The immersion took place at the dark limb at	17 10 2
The reappearance at the light limb at	18 27 44

" The immersion was perfectly instantaneous.

" In consequence of the strong light of the moon, which was nearly full, and the emersion taking place at the bright limb, I had no chance of perceiving the small star with the telescope employed, my equatoreal of 4 feet focal length and $3\frac{1}{8}$ inches aperture. The power used was 173. Longitude 23$^{\mathrm{s}}$·8 W.

" *Highbury, 8th July*, 1856."

Observations of the New Planet Isis, made with the Ten-feet Equatoreal and Ring Micrometer, at the Radcliffe Observatory, Oxford. By Mr. Norman Pogson.

(*Communicated by the President.*)

Greenwich M.T.		App. R.A.	App. N.P.D.	Logarithm of Par. Corr. × Δ. R.A.	Logarithm of Par. Corr. × Δ. N.P.D.	Star of Comp.
	h m s	h m s	° ′ ″			
May 28	12 33 44	16 7 36·07	...	+ 8·884	...	*a*
	13 3 0	16 7 34·73	105 30 50·7	9·084	−0·893	*a*
	13 37 19	16 7 33·14	105 30 54·7	9·238	·788	*a*
June 1	11 26 47	16 3 16·07	105 39 19·0	7·206	·898	*b*
	11 26 47	16 3 16·39	105 39 19·1	7·206	·898	*c*
2	13 1 36	16 2 7·68	105 41 48·2	9·197	·890	*c*
3	13 44 25	16 1 1·93	105 44 14·8	9·349	·881	*c*
6	12 23 59	15 57 57·87	105 51 41·5	9·080	·895	*c*
9	13 29 57	15 54 54·96	106 0 10·7	9·384	·879	*d*
	13 46 3	15 54 54·46	106 0 13·8	9·417	·874	*d*
10	12 8 7	15 54 0·38	106 2 59·0	+9·136	−0·893	*d*

All the observations have been duly corrected for refraction and motion. The adopted places of the comparison stars, for the epoch 1856·0, were as follows:—

	Mag.	R.A.	N.P.D.
		h m s	° ′ ″
a = A.Z. 297.78 = 29610 Lal.	8	16 7 42·92	105 31 11·7
b = A.Z. 205.73	9	16 4 39·43	105 38 28·8
c = A.Z. 205.70 = A.Z. 297.70	$9\frac{1}{4}$	16 0 31·22	105 42 19·8
d = A.Z. 205.66 = A.Z. 297.65 = 29175 Lal.	$8\frac{1}{2}$	15 55 22·60	105 59 55·4

Argelander's positions have been allowed double weight to those of Lalande. The right ascension of star *d* is one minute too small in Argelander's Zone 297.*

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* With the exception of *Flora*, *Isis* is the nearest to the sun, and has, therefore, the shortest periods of the known asteroids.

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